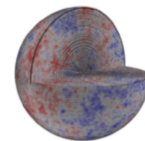




The terms Eisenstein & Hu Missed

A precise symbolic emulator of the linear matter power spectrum



Debating the potential of machine learning for astronomical surveys (#2)

27 November 2023

Deaglan Bartlett

Institut d'Astrophysique de Paris
CNRS & Sorbonne Université

deaglan.bartlett@iap.fr

Lukas Kammerer, Gabriel Kronberger, Bogdan Burlacu

Heuristic and Evolutionary Algorithms Laboratory, University of Applied Sciences Upper Austria

Benjamin D. Wandelt

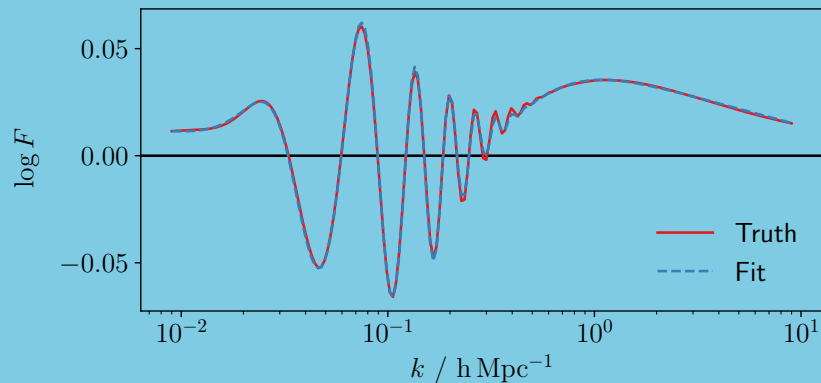
Institut d'Astrophysique de Paris & Center for Computational Astrophysics

Harry Desmond

Institute of Cosmology & Gravitation, University of Portsmouth

Pedro G. Ferreira, David Alonso, Matteo Zennaro

Astrophysics, University of Oxford



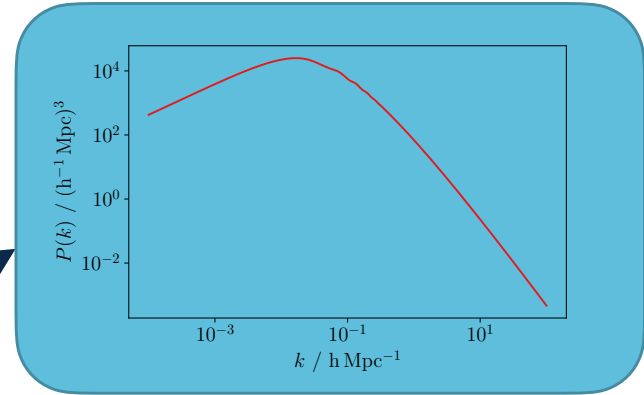
We don't always want to use a Boltzmann solver

Cosmological Parameters

- A_s
- Ω_m
- Ω_b
- h
- n_s

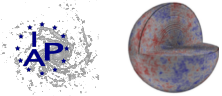
Boltzmann Solver

Power Spectra



Derived Parameters

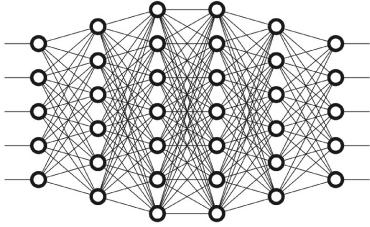
σ_8



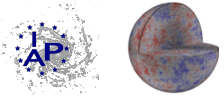
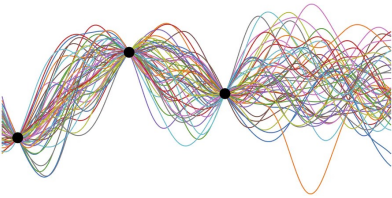
Many options exist instead of using Boltzmann solver

Numerical

Aricò et al. 2021;
Spurio Mancini et al. 2022
Euclid Collaboration et al. 2021



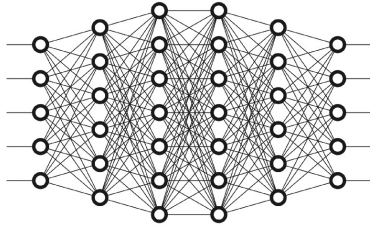
Mootoovaloo et al. 2022



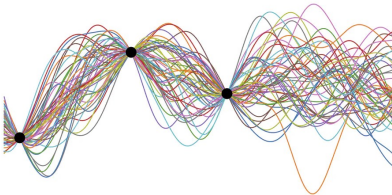
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Motoovaloo et al. 2022



Symbolic

Eisenstein & Hu 1998, 1998
BBKS 1986

The transfer function is written as a sum of the baryon and cold dark matter contributions at the drag epoch

$$T(k) = \frac{\Omega_b}{\Omega_0} T_b(k) + \frac{\Omega_c}{\Omega_0} T_c(k). \quad (8)$$

The CDM transfer function can be solved exactly in terms of hypergeometric functions that are more conveniently approximated by the following form:

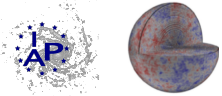
$$T_c \rightarrow \alpha_c \frac{\ln 1.8 \beta_c q}{14.2 q^2}, \quad (9)$$

$$q = \left(\frac{k}{\text{Mpc}^{-1}} \right) \Theta_{2.7}^2 (\Omega_0 h^2)^{-1} = \frac{k}{13.41 k_{\text{eq}}}, \quad (10)$$

where α_c and β_c are fitted by

$$\begin{aligned} \alpha_c &= a_1^{-\Omega_b/\Omega_0} a_2^{-(\Omega_b/\Omega_0)^3}, \\ a_1 &= (46.9 \Omega_0 h^2)^{0.670} [1 + (32.1 \Omega_0 h^2)^{-0.532}], \\ a_2 &= (12.0 \Omega_0 h^2)^{0.424} [1 + (45.0 \Omega_0 h^2)^{-0.582}], \\ \beta_c^{-1} &= 1 + b_1 [(\Omega_c/\Omega_0)^{b_2} - 1], \\ b_1 &= 0.944 [1 + (458 \Omega_0 h^2)^{-0.708}]^{-1}, \\ b_2 &= (0.395 \Omega_0 h^2)^{-0.0266}. \end{aligned} \quad (11)$$

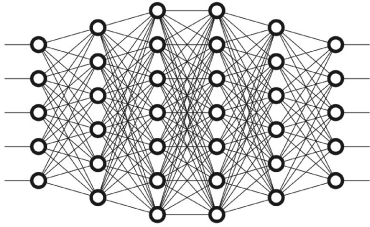
As $\Omega_b/\Omega_0 \rightarrow 0$, $\alpha_c \rightarrow 1$. Equation (9) shows the familiar $\ln(k)/k^2$ dependence of the small-scale CDM transfer function.



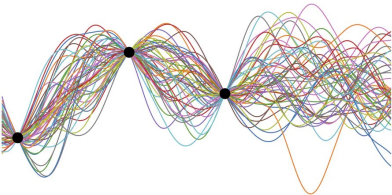
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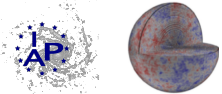
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- Interpretable
- Easily incorporate Physics
- Portable
- Longevity
- Fewer parameters



but ...

Not accurate enough



Find symbolic emulators with Symbolic Regression



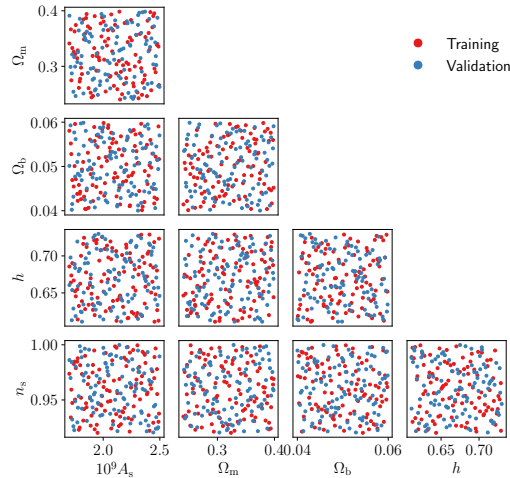
- Multi-Objective Genetic Programming
- Fast
- Memory efficient
- High performance on benchmarks
- C++ with python wrapper



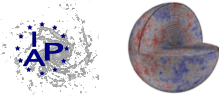
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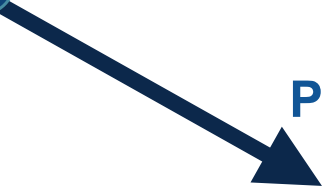
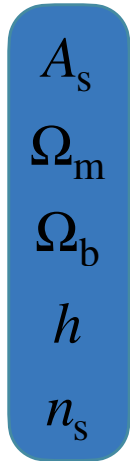


- 100 training, 100 validation cosmologies on latin hypercube (small number compared to e.g. neural networks)
- Truth computed with CAMB



We don't always want to use a Boltzmann solver

Cosmological Parameters

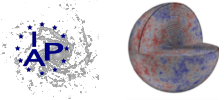


Derived Parameters



Task: Map $\sigma_8 \leftrightarrow A_s$ given other cosmological parameters

Normally: Run a Boltzmann solver with a guess of A_s and measure σ_8

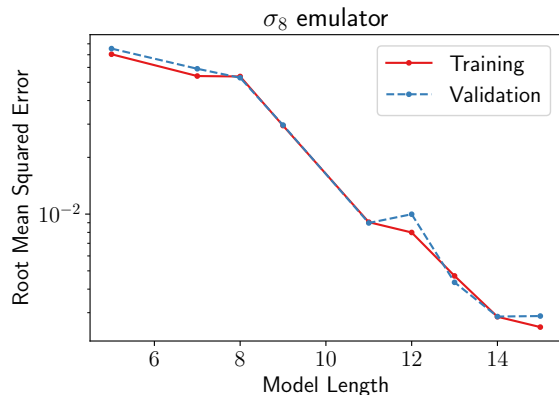


Task: Map $\sigma_8 \leftrightarrow A_s$ given other cosmological parameters

Normally: Run a Boltzmann solver with a guess of A_s and measure σ_8

Now:

$$\sigma_8 \approx (a_0 A_s + a_1 n_s) \left(a_2 \Omega_b + \log(a_3 \Omega_m) \right) \log(a_4 h) + a_5$$



- One simple line
- RMSE of 0.3% across range of cosmologies



We don't always want to use a Boltzmann solver

Cosmological Parameters

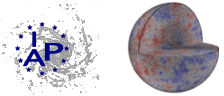
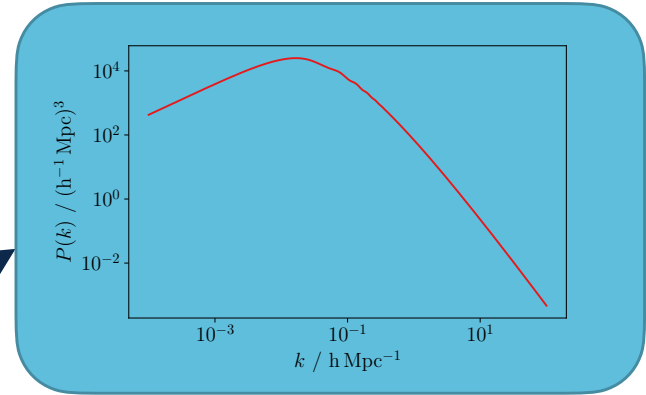
- A_s
- Ω_m
- Ω_b
- h
- n_s



Boltzmann Solver



Power Spectra



Eisenstein & Hu did the hard work! Fit the residuals

$$P(k, \theta) = P_{\text{EH}}(k, \theta)F(k, \theta)$$



Eisenstein & Hu did the hard work! Fit the residuals

$$P(k, \theta) = P_{\text{EH}}(k, \theta)F(k, \theta)$$

Physics!

- Baryonic acoustic oscillations
- Compton drag
- Velocity overshoot
- Baryon infall
- Adiabatic damping
- Silk damping
- Cold dark matter growth suppression



Eisenstein & Hu did the hard work! Fit the residuals

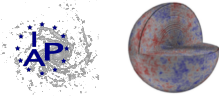
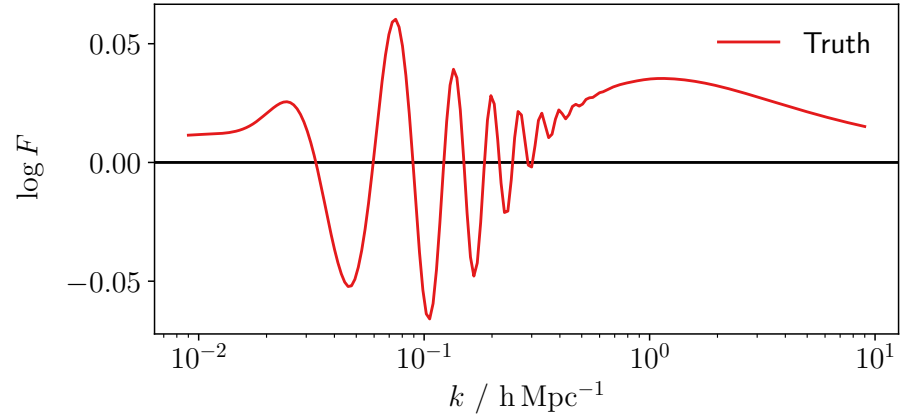
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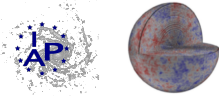
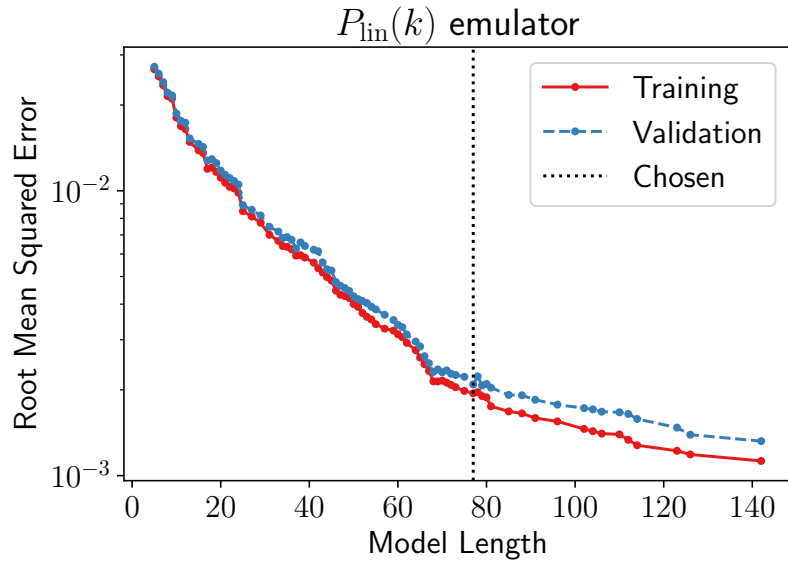
- Baryonic-acoustic oscillations
- Compton drag
- Velocity overshoot
- Baryon infall
- Adiabatic damping
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The rest

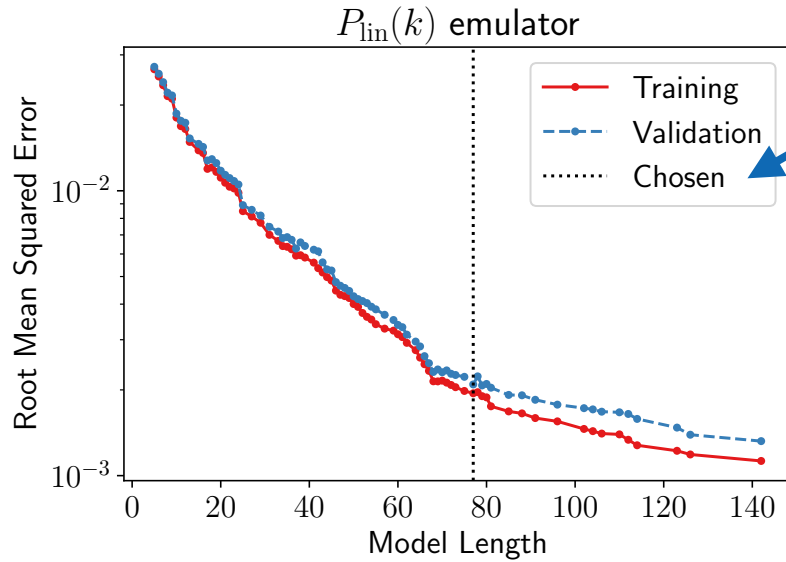
We will learn $\log F$



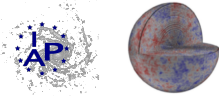
Percent-level fit achieved with (relatively) short expressions



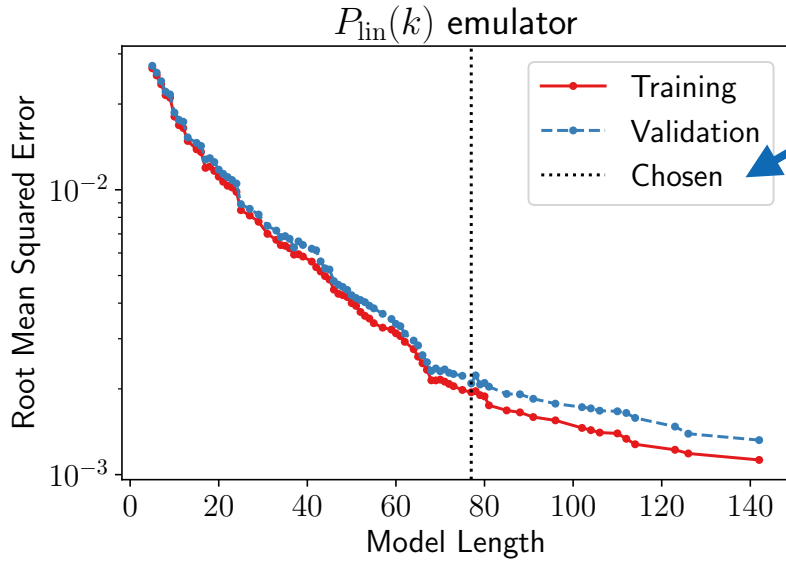
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Sufficient for % level accuracy



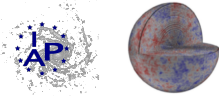
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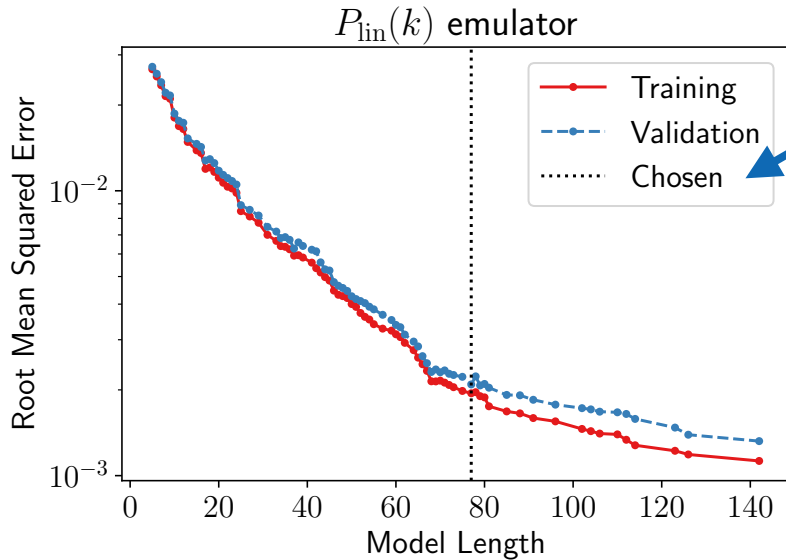
Sufficient for % level accuracy

Physically interpretable:

$$\sim \cos \left(b_{10} \Omega_m - \frac{b_{11} k}{\sqrt{b_{12} + \Omega_b^2}} \right)$$



Percent-level fit achieved with (relatively) short expressions

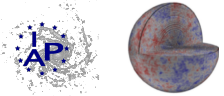


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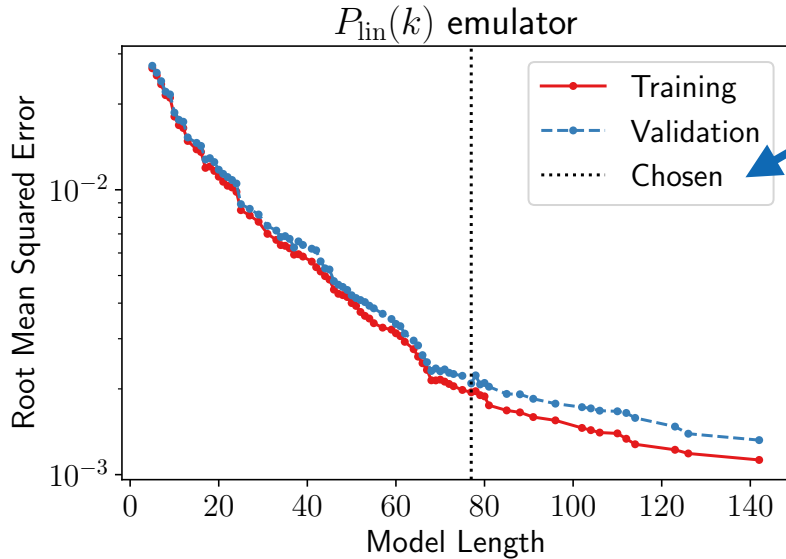
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BAO phase shift depends on Ω_m



Percent-level fit achieved with (relatively) short expressions



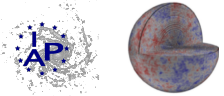
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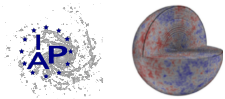
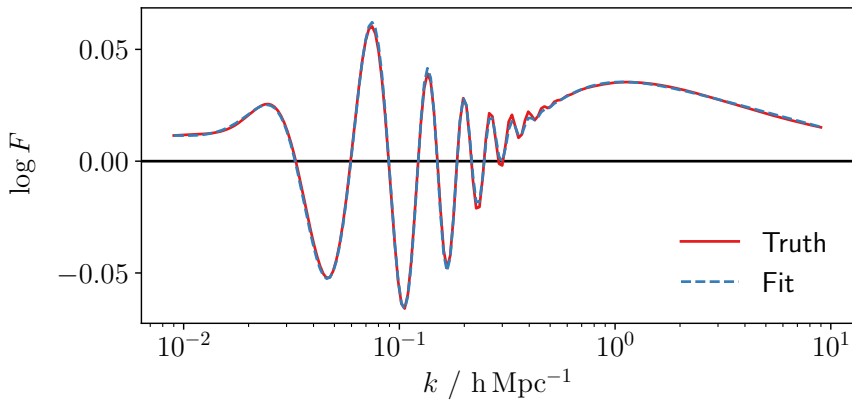
BAO phase shift depends on Ω_m

More BAO cycles in given k range for larger Ω_b



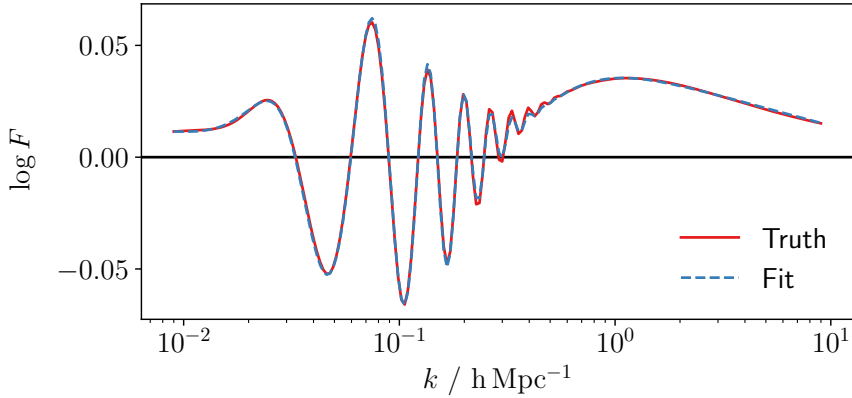
Symbolic emulator with sub-percent accuracy (RMSE 0.2%)

Fit almost perfect to the eye:

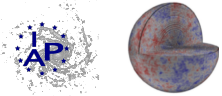
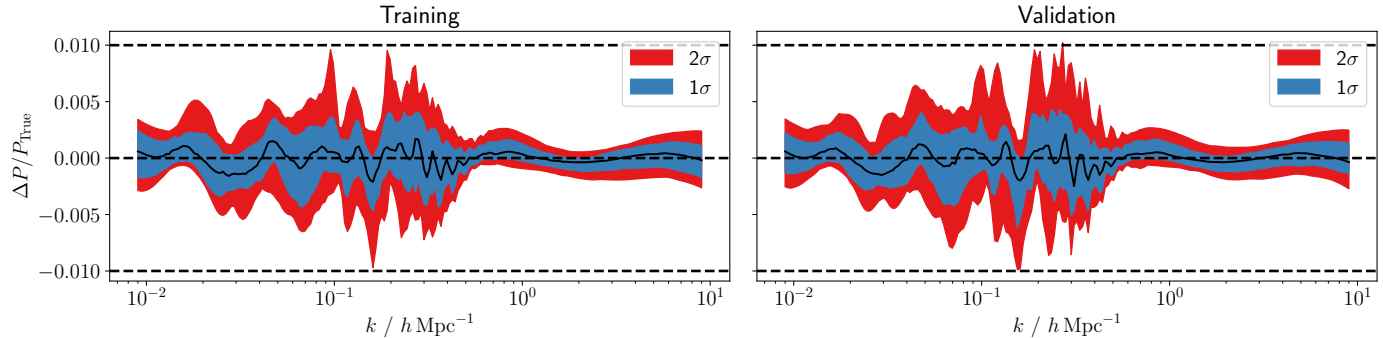


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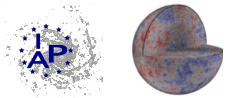
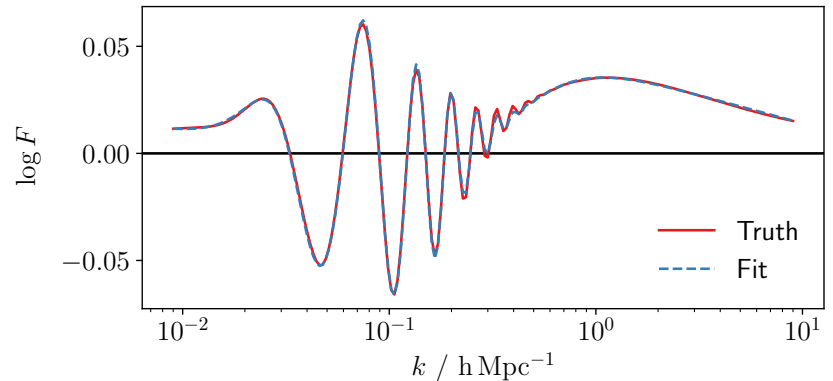


Sub-percent accuracy across training and validation:



Conclusions

- Symbolic emulator for linear $P(k)$ with sub-percent accuracy
- One line approximation for $\sigma_8 \leftrightarrow A_s$ with RMSE of 0.3%
- Do I *really* need a neural network?
 - Portability
 - Interpretability
 - Longevity
 - Fewer parameters



Extra Slides



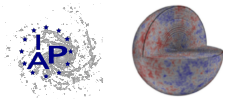
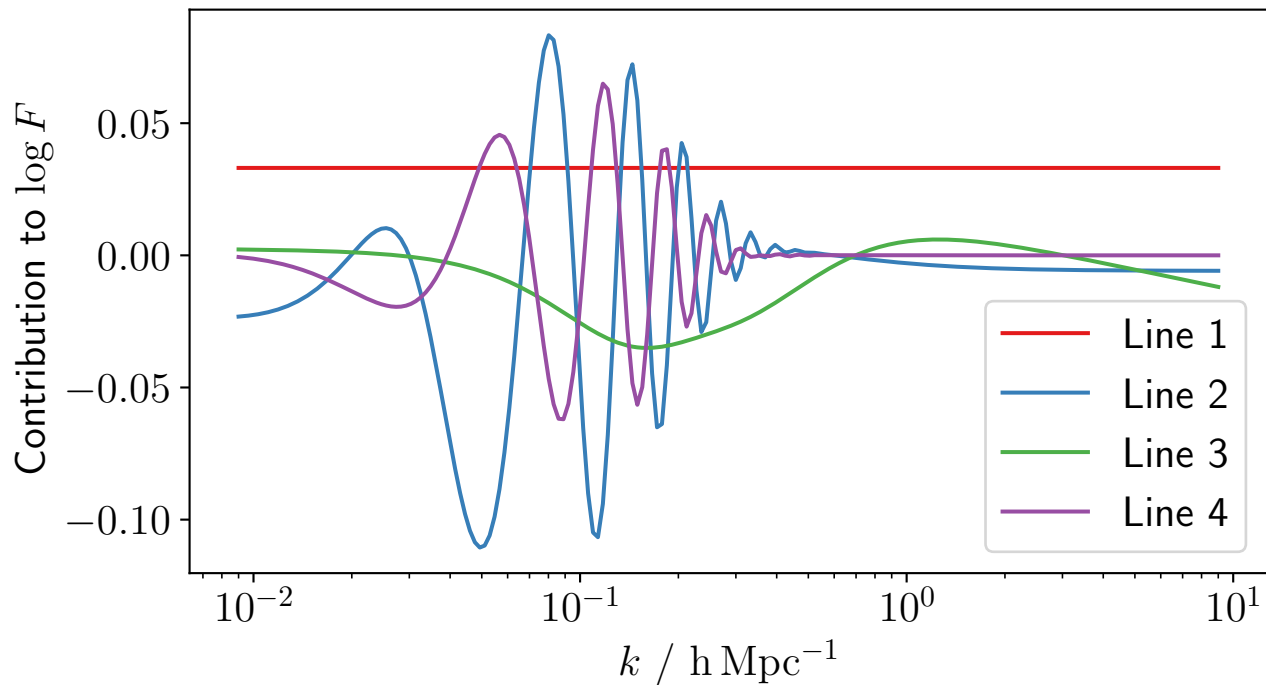
Symbolic emulator for non-linear power spectrum

$$\log F \approx b_0 h - b_{26}$$

$$\begin{aligned}
 & + \left(\frac{b_1 \Omega_b}{\sqrt{h^2 + b_2}} \right)^{b_3 \Omega_m} \left[\frac{b_4 k - \Omega_b}{\sqrt{b_5 + (\Omega_b - b_6 k)^2}} b_7 (b_8 k)^{-b_9 k} \cos \left(b_{10} \Omega_m - \frac{b_{11} k}{\sqrt{b_{12} + \Omega_b^2}} \right) - b_{13} \left(\frac{b_{14} k}{\sqrt{1 + b_{15} k^2}} - \Omega_m \right) \cos \left(\frac{b_{16} h}{\sqrt{1 + b_{17} k^2}} \right) \right] \\
 & + b_{18} (b_{19} \Omega_m + b_{20} h - \log(b_{21} k) + (b_{22} k)^{-b_{23} k}) \cos \left(\frac{b_{24}}{\sqrt{1 + b_{25} k^2}} \right) \\
 & + (b_{27} k)^{-b_{28} k} \left(b_{29} k - \frac{b_{30} \log(b_{31} k)}{\sqrt{b_{32} + (\Omega_m - b_{33} h)^2}} \right) \cos \left(b_{34} \Omega_m - \frac{b_{35} k}{\sqrt{b_{36} + \Omega_b^2}} \right)
 \end{aligned}$$



Contributions to $P(k)$ emulator



Residuals for Planck 2018 Cosmology

